








# Intermediate-mass black holes from stellar mergers in young star clusters

Ugo N. Di Carlo <sup>1,2,3,4</sup>★ Michela Mapelli <sup>1,2,3</sup> Mario Pasquato,<sup>5</sup> Sara Rastello <sup>1,2</sup>  
Alessandro Ballone <sup>1,2,3</sup> Marco Dall’Amico,<sup>1,2</sup> Nicola Giacobbo <sup>1,6</sup> Giuliano Iorio <sup>1,2,3</sup>  
Mario Spera <sup>1,2,7</sup> Stefano Tornamenti<sup>1,2,3</sup> and Francesco Haardt<sup>4</sup>

<sup>1</sup>Dipartimento di Fisica e Astronomia ‘G. Galilei’, University of Padova, Vicolo dell’Osservatorio 3, I-35122 Padova, Italy

<sup>2</sup>INFN, Sezione di Padova, Via Marzolo 8, I-35131 Padova, Italy

<sup>3</sup>INAF-Osservatorio Astronomico di Padova, Vicolo dell’Osservatorio 5, I-35122 Padova, Italy

<sup>4</sup>Dipartimento di Scienza e Alta Tecnologia, University of Insubria, Via Valleggio 11, I-22100 Como, Italy

<sup>5</sup>Center for Astro, Particle and Planetary Physics (CAP<sup>3</sup>), New York University Abu Dhabi, Abu Dhabi, UAE

<sup>6</sup>School of Physics and Astronomy, Institute for Gravitational Wave Astronomy, University of Birmingham, Birmingham B15 2TT, UK

<sup>7</sup>Scuola Internazionale Superiore di Studi Avanzati (SISSA), Via Bonomea 265, I-34136 Trieste, Italy

Accepted 2021 August 16. Received 2021 August 6; in original form 2021 May 5

## ABSTRACT

Intermediate-mass black holes (IMBHs) in the mass range  $10^2$ – $10^5 M_{\odot}$  bridge the gap between stellar black holes (BHs) and supermassive BHs. Here, we investigate the possibility that IMBHs form in young star clusters via runaway collisions and BH mergers. We analyse  $10^4$  simulations of dense young star clusters, featuring up-to-date stellar wind models and prescriptions for core collapse and (pulsational) pair instability. In our simulations, only nine IMBHs out of 218 form via binary BH mergers, with a mass  $\sim 100$ – $140 M_{\odot}$ . This channel is strongly suppressed by the low escape velocity of our star clusters. In contrast, IMBHs with masses up to  $\sim 438 M_{\odot}$  efficiently form via runaway stellar collisions, especially at low metallicity. Up to  $\sim 0.2$  per cent of all the simulated BHs are IMBHs, depending on progenitor’s metallicity. The runaway formation channel is strongly suppressed in metal-rich ( $Z = 0.02$ ) star clusters, because of stellar winds. IMBHs are extremely efficient in pairing with other BHs:  $\sim 70$  per cent of them are members of a binary BH at the end of the simulations. However, we do not find any IMBH–BH merger. More massive star clusters are more efficient in forming IMBHs:  $\sim 8$  per cent ( $\sim 1$  per cent) of the simulated clusters with initial mass  $10^4$ – $3 \times 10^4 M_{\odot}$  ( $10^3$ – $5 \times 10^3 M_{\odot}$ ) host at least one IMBH.

**Key words:** black hole physics – gravitational waves – methods: numerical – binaries: general – stars: kinematics and dynamics – galaxies: star clusters: general.

## 1 INTRODUCTION

Intermediate-mass black holes (IMBHs), with mass in the range  $10^2$ – $10^5 M_{\odot}$ , bridge the gap between stellar-sized black holes (BHs) and supermassive BHs. IMBHs might be the building blocks for supermassive BHs in the early Universe (Madau & Rees 2001), as suggested by the discovery of quasars at redshift as high as  $\sim 6$  (e.g. Fan et al. 2001). Currently, we lack unambiguous evidence of IMBH existence from electromagnetic observations, but a number of potential candidates were proposed in the last decades. Several globular clusters, including 47 Tucanae (Kızıltan, Baumgardt & Loeb 2017) and G1 (Gebhardt, Rich & Ho 2005), have been claimed to host IMBHs, based on indirect dynamical pieces of evidence (Gebhardt, Rich & Ho 2002; Gerssen et al. 2002; Anderson & van der Marel 2010; van der Marel & Anderson 2010; Lützgendorf et al. 2011, 2013; Miller-Jones et al. 2012; Nyland et al. 2012; Strader et al. 2012; Lanzoni et al. 2013; Baumgardt 2017; Perera et al. 2017; Lin et al. 2018; Tremou et al. 2018; Baumgardt et al. 2019; Mann et al. 2019; Zocchi, Gieles & Hénault-Brunet 2019). The hyperluminous

X-ray source HLX-1 in the galaxy ESO 243-49 (Farrell et al. 2009; Godet et al. 2014) is possibly the strongest IMBH candidate among ultra-luminous X-ray sources (e.g. Kaaret et al. 2001; Matsumoto et al. 2001; Strohmayer & Mushotzky 2003; Miller & Colbert 2004; Swartz et al. 2004; van der Marel 2004; Mapelli et al. 2010; Feng & Soria 2011; Sutton et al. 2012; Mezcua et al. 2015; Kaaret, Feng & Roberts 2017; Lin et al. 2020). Several BHs with mass  $\sim 10^5 M_{\odot}$  might lurk in the nuclei of dwarf galaxies (Filippenko & Sargent 1989; Filippenko & Ho 2003; Barth et al. 2004; Greene & Ho 2004, 2007; Seth et al. 2010; Dong et al. 2012; Kormendy & Ho 2013; Baldassare et al. 2015; Reines & Volonteri 2015; Mezcua et al. 2016, 2018; Reines et al. 2020). Finally, some candidates with mass  $\sim 10^4$ – $10^5 M_{\odot}$  may lie in high-velocity compact clouds (Tanaka 2018; Takekawa et al. 2019, 2020). We point to Mezcua (2017) and Greene, Strader & Ho (2020) for two recent reviews.

The gravitational wave event GW190521 (Abbott et al. 2020a, b), associated with the merger of a binary BH (BBH) with a total mass of  $\sim 150 M_{\odot}$ , recently provided strong evidence for IMBHs and paved the ground for further observations with current (Aasi et al. 2015; Acernese et al. 2015) and future gravitational wave detectors (Punturo et al. 2010; Amaro-Seoane et al. 2017; Kalogera et al. 2019; Reitze et al. 2019; Maggiore et al. 2020). Moreover,

\* E-mail: ugo.dc@hotmail.it

GW190521 confirms the merger formation scenario for IMBHs, where smaller BHs, likely of stellar origin, repeatedly merge in a dense environment, such as a star cluster (SC, Miller & Hamilton 2002; McKernan et al. 2012; Giersz et al. 2015; Fishbach, Holz & Farr 2017; Gerosa & Berti 2017, 2019; Rodriguez et al. 2019; Antonini, Gieles & Gualandris 2019; Arca Sedda et al. 2020; Fragione, Loeb & Rasio 2020; Mapelli et al. 2020a, 2021; Tagawa et al. 2021a, b). Other possible formation scenarios are the collapse of very massive ( $>230 M_{\odot}$ ) metal-poor stars (Bond, Arnett & Carr 1984; Madau & Rees 2001; Heger & Woosley 2002; Schneider et al. 2002; Spera & Mapelli 2017; Tanikawa et al. 2021), primordial IMBHs from gravitational instabilities in the early Universe (Kawaguchi et al. 2008; Carr, Kühnel & Sandstad 2016; Raccanelli et al. 2016; Sasaki et al. 2016; Scelfo et al. 2018; Clesse & Garcia-Bellido 2020; De Luca et al. 2021) and runaway collisions of stars in SCs (Colgate 1967; Sanders 1970; Portegies Zwart & McMillan 2002; Gürkan, Freitag & Rasio 2004; Portegies Zwart et al. 2004; Freitag, Rasio & Baumgardt 2006a; Giersz et al. 2015; Mapelli 2016; Di Carlo et al. 2019; Chon & Omukai 2020; Kremer et al. 2020; Das et al. 2021; Rizzuto et al. 2021).

Different formation channels may be at work in different environments (Volonteri 2010), with SCs representing the ideal place both for repeated mergers of stellar-mass BHs (e.g. Miller & Hamilton 2002) and for the formation of very massive stars through runaway collisions in the early stages of the cluster’s evolution (e.g. Fujii & Portegies Zwart 2013). The internal dynamics of SCs is expected to play a key role in modulating the efficiency of the BBH merger scenario, with core collapse being closely tied to increased dynamical interactions (Spitzer 1987; Goodman & Hut 1989) and to the dynamical hardening of binary systems (Sugimoto & Bettwieser 1983; Goodman & Hut 1993; Trenti et al. 2007; Hurley & Shara 2012; Beccari et al. 2019; Di Cintio et al. 2021). The main advantage of the repeated merger mechanism is that it works for the entire lifetime of an SC: the IMBH can assemble and grow as long as the SC survives (Giersz et al. 2015). However, a merger chain can come to an abrupt end if the BBH merger product is ejected from the SC due to gravitational-wave recoil (Madau & Quataert 2004; Merritt et al. 2004; Holley-Bockelmann et al. 2008). The magnitude of such kick can largely exceed the escape velocity of the host SC (Favata, Hughes & Holz 2004; Lousto & Zlochower 2011), thus the hierarchical build-up of an IMBH via repeated BBH mergers is more likely to take place in the most massive SCs, like nuclear SCs, because of their large escape velocity (Antonini et al. 2019; Fragione et al. 2020; Mapelli et al. 2021).

The runaway collision mechanism, on the other hand, can be effective even in less massive SCs, like young SCs. The most massive stars in a young SC are more likely to undergo collisions than lighter stars, because dynamical friction quickly brings them to the core (Gaburov, Lombardi & Portegies Zwart 2008). The high central density of young SCs, which is further enhanced in the first Myr due to gravothermal collapse, greatly favours stellar collisions (Freitag, Gürkan & Rasio 2006b; Portegies Zwart, McMillan & Gieles 2010), which help to build up very massive stars that may collapse into IMBHs. The main issue is that very massive stars lose a conspicuous portion of their mass through stellar winds (e.g. Vink et al. 2011). The IMBH can form only if the star preserves enough mass, so this mechanism is more efficient at lower metallicity, where stellar winds are less powerful (Mapelli 2016). Moreover, all the collisions need to take place before the massive stars in the SC turn into BHs, which happens in a few Myr.

Despite being less massive than globular clusters and nuclear SCs, young SCs make up the vast majority of SCs in the Universe

**Table 1.** Initial conditions.

Set	$Z$	$N_{\text{sim}}$	$r_{\text{h}}$	$D$	ref.
D2019HF	0.002	2000	M12	2.3	D19
D2019LF	0.002	2000	M12	1.6	D19
D2020A	0.02	1000	M12	1.6	D20
	0.002	1000	M12	1.6	D20
	0.0002	1000	M12	1.6	D20
D2020B	0.02	1000	1.5 pc	1.6	D20
	0.002	1000	1.5 pc	1.6	D20
	0.0002	1000	1.5 pc	1.6	D20

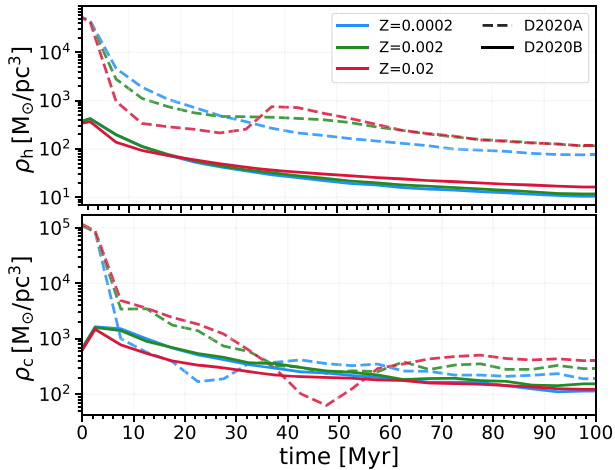
*Notes.* Column 1: name of the simulation set; column 2: metallicity  $Z$ ; column 3: number of runs performed per each set; column 4: half-mass radius  $r_{\text{h}}$ . M12 indicates that half-mass radii have been drawn according to Marks & Kroupa (2012). Column 5: fractal dimension ( $D$ ); column 6: reference for each simulation set. D19 and D20 correspond to Di Carlo et al. (2019) and Di Carlo et al. (2020a), respectively.

(Kroupa & Boily 2002), and their cumulative contribution to IMBH statistics may thus be significant. In this paper, we study the demography of IMBHs in young SCs through  $N$ -body simulations, including up-to-date stellar wind models and a treatment for (pulsational) pair instability in massive and very massive stars. We present a set of  $10^4$  state-of-the-art direct  $N$ -body simulations with fractal initial conditions (to mimic the clumpiness of star-forming regions, e.g. Goodwin & Whitworth 2004) and with a large initial binary fraction ( $f_{\text{bin}} = 1$  for stars with mass  $>5 M_{\odot}$ ).

## 2 METHODS

The simulations discussed in this paper were performed using the same codes and methodology described in Di Carlo et al. (2019). We use the direct summation  $N$ -Body code NBODY6+ + GPU (Wang et al. 2015), which we coupled with the population synthesis code MOBSE (Mapelli et al. 2017; Giacobbo, Mapelli & Spera 2018). MOBSE includes up-to-date prescriptions for massive stellar winds (Giacobbo & Mapelli 2018), electron-capture supernovae (Giacobbo & Mapelli 2019), core-collapse supernovae (Fryer et al. 2012), natal kicks (Giacobbo & Mapelli 2020), and (pulsational) pair-instability supernovae (Mapelli et al. 2020b). Orbital decay induced by gravitational wave emission is treated as described in Peters (1964). If a star merges with a BH or with a neutron star, MOBSE assumes that the entire mass of the star is lost by the system and not accreted by the compact object.

In this work, we have analysed  $10^4$  simulations of young SCs; 4000 of them are the simulations presented in Di Carlo et al. (2019), while the remaining 6000 are the ones discussed in Di Carlo et al. (2020a). The initial conditions of the simulations are summarized in Table 1. Young SCs are asymmetric, clumpy systems. Thus, we model them with fractal initial conditions (Küpper et al. 2011), to mimic the clumpiness of star-forming regions (Goodwin & Whitworth 2004). The level of fractality is decided by the parameter  $D$ , where  $D = 3$  means homogeneous distribution of stars. In this work, we assume  $D = 1.6$  in most of our simulations (8000 runs). We also probe a lower degree of fractality ( $D = 2.3$ ) in the remaining 2000 runs. The total initial mass  $M_{\text{SC}}$  of each SC ranges from  $10^3$  to  $3 \times 10^4 M_{\odot}$  and it is drawn from a distribution  $dN/dM_{\text{SC}} \propto M_{\text{SC}}^{-2}$ , as the SC initial mass function described in Lada & Lada (2003). Thus, the mass distribution of our simulated SCs mimics the mass distribution of SCs in Milky Way-like galaxies. We choose the initial half-mass radius  $r_{\text{h}}$  according to Marks & Kroupa (2012) in 7000 simulations,



**Figure 1.** Evolution of the median of the mass density at the half-mass radius ( $\rho_h$ , top panel) and at the core radius ( $\rho_c$ , bottom panel) for the SCs in sets D2020A ( $r_h$  from Marks & Kroupa 2012, dashed lines) and D2020B ( $r_h = 1.5$  pc, solid lines) at different metallicities.

and we adopt a fixed value  $r_h = 1.5$  pc for the remaining 3000 simulations. These different choices of  $r_h$  have a strong impact on the density of the SCs, as shown in Fig. 1.

We generated the initial conditions with MCLUSTER (Küpper et al. 2011). The initial masses of the stars follow a Kroupa (2001) initial mass function, with minimum mass  $0.1 M_{\odot}$  and maximum mass  $150 M_{\odot}$ . We set an initial total binary fraction of  $f_{\text{bin}} = 0.4$ . Since MCLUSTER pairs up the most massive primary stars first, according to a distribution  $\mathcal{P}(q) \propto q^{-0.1}$  (where  $q = m_2/m_1 \in [0.1, 1]$ , Sana et al. 2012), all the stars with mass  $m \geq 5 M_{\odot}$  are members of binary systems, while stars with mass  $m < 5 M_{\odot}$  are randomly paired until a total binary fraction  $f_{\text{bin}} = 0.4$  is reached. This procedure results in a mass-dependent initial binary fraction consistent with the multiplicity properties of O/B-type stars (Sana et al. 2012; Moe & Di Stefano 2017). The orbital periods and eccentricities are also drawn from the distributions described in Sana et al. (2012).<sup>1</sup>

We simulate SCs with three different metallicities:  $Z = 0.0002$ ,  $0.002$ , and  $0.02$  (approximately  $1/100$ ,  $1/10$ , and  $1 Z_{\odot}$ , assuming  $Z_{\odot} = 0.02$  from Anders & Grevesse 1989). Each SC is simulated for 100 Myr in a solar neighbourhood-like static external tidal field (Wang et al. 2016). At 100 Myr, we extract the BBHs which are still present in our simulations and we integrate their semimajor axis and eccentricity evolution by gravitational wave emission up to a Hubble time (13.6 Gyr), according to Peters (1964).

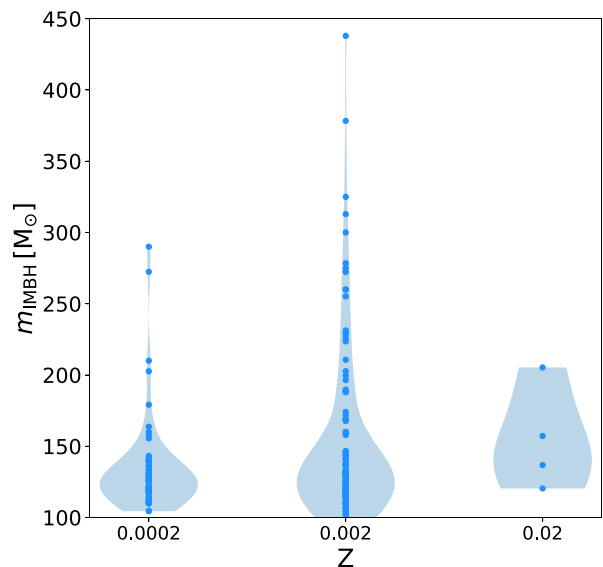
The main differences between the simulations in Di Carlo et al. (2019) and Di Carlo et al. (2020a) are the efficiency of common envelope ejection ( $\alpha = 3$  in Di Carlo et al. 2019 and  $\alpha = 5$  in Di Carlo et al. 2020a), and the chosen model of core-collapse supernova (the rapid and the delayed models from Fryer et al. 2012 are adopted in Di Carlo et al. 2019 and in Di Carlo et al. 2020a, respectively). The population of IMBHs is not strongly affected by these differences. We divide the simulations in four different sets whose names and properties are shown in Table 1. We refer to Di Carlo et al. (2019) and Di Carlo et al. (2020a) for further details on the code and on the simulations.

<sup>1</sup>The same binary properties could have been obtained using the binary eigen-evolution and mass-feeding method described in Belloni et al. (2018).

**Table 2.** Percentage of SCs that form 0, 1, and 2 IMBHs per different metallicity, initial SC mass, and simulation set.

Set	0 IMBHs	1 IMBH	2 IMBHs
All	97.97 per cent	1.97 per cent	0.06 per cent
$Z = 0.02$	99.85 per cent	0.15 per cent	0 per cent
$Z = 0.002$	97.81 per cent	2.17 per cent	0.02 per cent
$Z = 0.0002$	96.5 per cent	3.25 per cent	0.25 per cent
$10^3 \leq M_{\text{SC}}/M_{\odot} < 5 \times 10^3$	99.01 per cent	0.99 per cent	0 per cent
$5 \times 10^3 \leq M_{\text{SC}}/M_{\odot} < 10^4$	95.73 per cent	4.12 per cent	0.15 per cent
$10^4 \leq M_{\text{SC}}/M_{\odot} \leq 3 \times 10^4$	92.13 per cent	7.42 per cent	0.45 per cent
D2019HF	96.90 per cent	3.05 per cent	0.05 per cent
D2019LF	98.2 per cent	1.80 per cent	0.0 per cent
D2020A	98.03 per cent	1.94 per cent	0.03 per cent
D2020B	98.47 per cent	1.40 per cent	0.13 per cent

*Notes.* Column 1: Simulation set; column 2: percentage of SCs that form no IMBHs; column 3: percentage of SCs that form one IMBH; column 4: percentage of SCs that form two IMBHs.



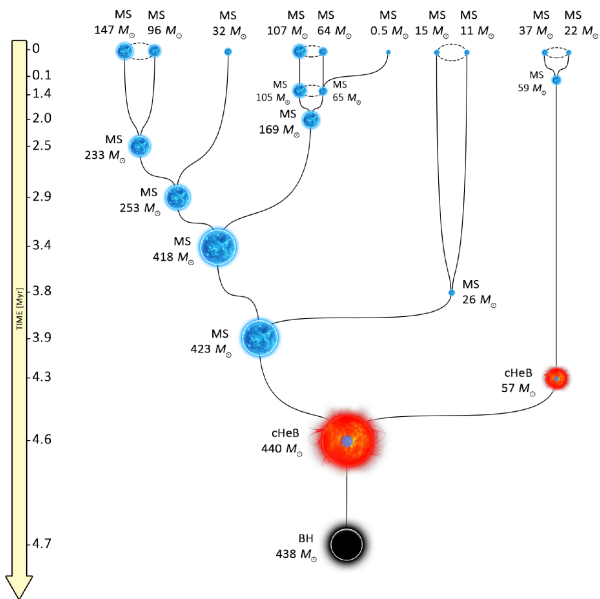
**Figure 2.** Distributions of IMBH masses for each metallicity. The blue filled circles mark the values of the masses, while the horizontal extent of each light blue region (violin plot) is proportional to the number of IMBHs at a given mass value.

## 3 RESULTS

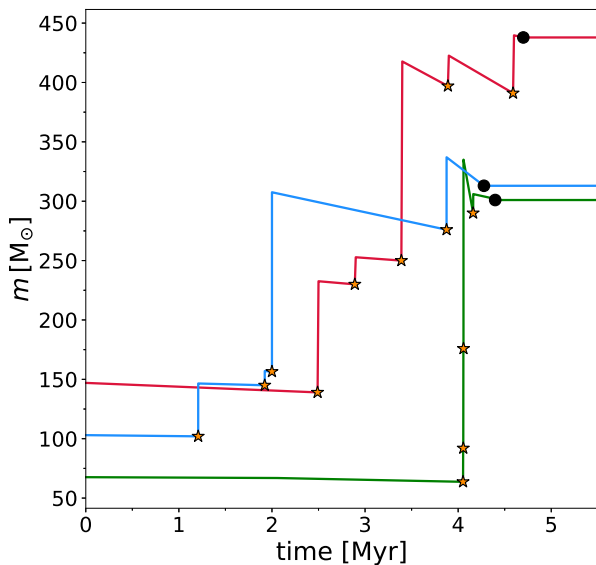
### 3.1 Formation of IMBHs

We find a total of 218 IMBHs, which make up only  $\sim 0.2$  per cent of all the BHs formed in our simulations. The formation of IMBHs is a rare event in our simulations: as shown in Table 2, only 1.97 per cent of all the simulated SCs form one IMBH, and only 0.06 per cent of them form two IMBHs. The vast majority of the simulated IMBHs (209) form via the runaway collision mechanism, while only nine IMBHs form via BBH merger. This is expected because the low escape velocity ( $\leq 10 \text{ km s}^{-1}$ ) of our simulated SCs makes the BBH merger scenario less efficient.

Fig. 2 shows the masses of all the formed IMBHs for each metallicity. IMBHs form with mass up to  $\sim 438 M_{\odot}$ , but  $\sim 78$  per cent of all the formed IMBHs has a mass between 100 and  $150 M_{\odot}$ . Less massive IMBHs are more likely to form. A schematic formation history of the most massive IMBH is shown in Fig. 3; this IMBH forms via runaway collision: a total of 10 stars participate to the

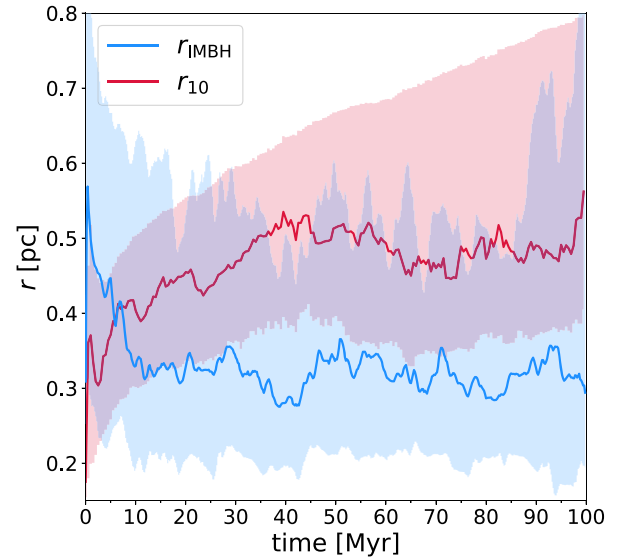


**Figure 3.** Formation history of the most massive IMBH found in our simulations. Main-sequence stars (with label MS) are represented as blue stars; core helium burning stars (label cHeB) are visualized as red stars with a blue core; black holes (label BH) are shown as black circles. The mass of each object is shown next to them. The time axis and the size of the objects are not to scale. The formation of this IMBH takes place in an SC with metallicity  $Z = 0.002$ .



**Figure 4.** Mass evolution of three stars which end up forming an IMBH. The yellow stars mark stellar mergers. The black filled circles indicate the moment when the IMBH form.

formation of a very massive star which promptly undergoes direct collapse to form the IMBH. All the other IMBHs which form via runaway collision follow a similar pattern, involving multiple stellar mergers in the first few Myr of our simulations. The mass evolution of three stars which end up forming an IMBH is shown in Fig. 4. We see that mass growth via stellar mergers overcomes stellar wind mass-loss. The formation of the IMBHs via runaway collision always occurs within the first  $\sim 5$  Myr from the beginning of the simulation.



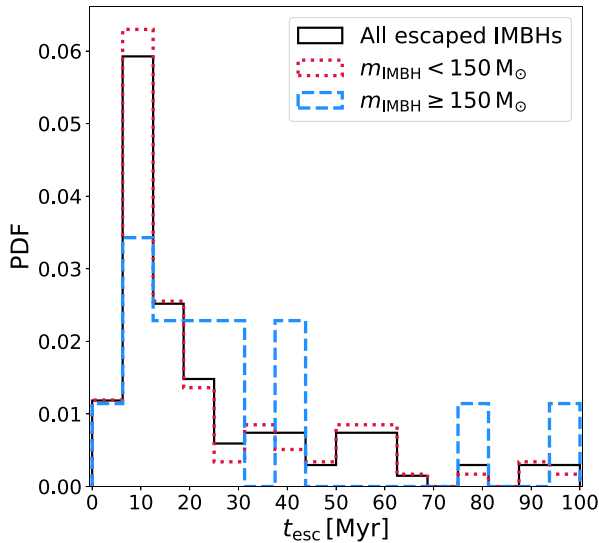
**Figure 5.** Evolution of the distance of an IMBH from the centre of mass of its host SC ( $r_{\text{IMBH}}$ ). The blue line shows the median of the values of  $r_{\text{IMBH}}$ , while the red line is the median of the 10 percent Lagrangian radii ( $r_{10}$ ) of the host SCs. The filled areas represent the 25th–75th percentile confidence intervals. Only IMBHs which do not escape from their host SC are shown in this figure. A simple moving average over five time-steps has been performed to filter out statistical fluctuations.

As expected, IMBH formation is much less efficient at higher metallicity because stellar winds are more powerful; only four IMBHs form at solar metallicity. The fraction of IMBHs with respect to the total number of BHs per each metallicity is  $\sim 0.4$  per cent,  $\sim 0.2$  per cent, and  $\sim 0.02$  per cent for  $Z = 0.0002$ ,  $0.002$ , and  $0.02$ , respectively. Table 2 shows that the percentage of simulated SCs which form at least one IMBH grows as the metallicity decreases, going from 0.15 per cent at  $Z = 0.02$  to 3.5 per cent at  $Z = 0.0002$ . We obtained a larger absolute number of IMBHs at metallicity  $Z = 0.002$  (Fig. 2), only because we ran three times more simulations with  $Z = 0.002$  than with either  $Z = 0.02$  or  $Z = 0.0002$  (Table 1). We performed several statistical tests to check if the fact that the most massive IMBHs form at  $Z = 0.002$  is statistically significant or is just another effect of the larger number of simulation at this metallicity. Our tests suggest that intermediate-metallicity SCs (with  $Z = 0.002 = 0.1 Z_{\odot}$ ) tend to form more massive IMBHs than metal-poor SCs (with  $Z = 0.0002 = 0.01 Z_{\odot}$ ), even if there is no conclusive evidence for this result (see Appendix A for details). This possible difference could be interpreted as the interplay between relatively inefficient mass-loss and large maximum stellar radii (Chen et al. 2015): at  $Z = 0.002$  stellar winds are still rather inefficient (as for lower metallicity), but massive stars can develop very large radii (in the terminal-age main sequence and the giant phase), enhancing the probability of collisions with other stars.

### 3.2 Distance from centre of mass

Fig. 5 shows the evolution of the median distance of the IMBHs with respect to the centre of mass of the host SCs. In order to undergo a sufficient number of stellar mergers, we expect the progenitor stars of IMBHs to lie in the densest central regions of the SC. In our simulations, we find that this is not completely the case: IMBH progenitors (i.e. before  $t \sim 5$  Myr) tend to lie outside the 10 percent Lagrangian radius. This is likely a consequence of the fractal initial





**Figure 6.** Probability distribution function of the escape time  $t_{\text{esc}}$  of IMBHs. All the ejected IMBHs are shown by the solid black line. IMBHs with mass  $m_{\text{IMBH}} < 150 M_{\odot}$  are shown by the dotted red line, while IMBHs with mass  $m_{\text{IMBH}} \geq 150 M_{\odot}$  are shown by the dashed blue line.

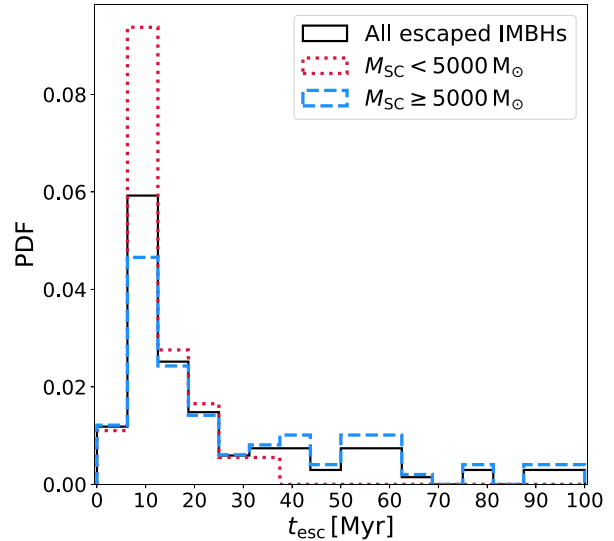
conditions, because the IMBH formation may take place in a clump far away from the centre of mass of the SC, before the mergers between the clumps take place. After the IMBHs form, they tend to rapidly sink towards the centre of the SC and to remain there until the end of the simulation.

### 3.3 Escaped IMBHs

Fig. 5 shows only IMBHs which do not escape from their host SC. In our simulations, particles escape from the SC when their distance from the centre of mass becomes larger than twice the tidal radius. We find that  $\sim 54$  per cent of all the formed IMBHs are ejected or evaporate from the SC in the first 100 Myr. The percentage of IMBHs with mass  $m_{\text{IMBH}} < 150 M_{\odot}$  that escape from their host SC is  $\sim 57.0$  per cent, and only  $\sim 31.8$  per cent for IMBHs with  $m_{\text{IMBH}} \geq 150 M_{\odot}$ . This means that the most massive IMBHs are more likely to be retained inside their host SC.

Fig. 6 shows the distribution of the escape times  $t_{\text{esc}}$  of ejected IMBHs. All IMBHs and especially the less massive ones tend to be ejected within the first 25 Myr, with a strong peak between 5 and 15 Myr. More massive IMBHs tend to escape at later times with respect to the less massive ones. Fig. 7 distinguishes between IMBHs that formed in SCs with initial mass smaller and larger than  $5 \times 10^3 M_{\odot}$ . Both distributions peak around 10 Myr, but all the ejected IMBHs which form in the less massive SCs escape within the first 40 Myr, while IMBHs born in more massive SCs may escape at later times.

We find that the percentages of escaped IMBHs which form in SCs with initial mass  $M_{\text{SC}} < 5000 M_{\odot}$  and  $M_{\text{SC}} \geq 5000 M_{\odot}$  are 37.7 per cent and 59.8 per cent, respectively. This means that IMBHs are more likely to be ejected from more massive SCs. This may seem surprising, because we expect more massive SCs to retain more IMBHs due to their larger escape velocities. Massive SCs, however, have denser cores where dynamical interactions that may eject IMBHs are more frequent.



**Figure 7.** Probability distribution function of the escape time  $t_{\text{esc}}$  of the IMBHs. All the ejected IMBHs are shown by the solid black line. IMBHs formed in SCs with mass  $M_{\text{SC}} < 5000 M_{\odot}$  are shown by the dotted red line, while IMBHs formed in SCs with mass  $M_{\text{SC}} \geq 5000 M_{\odot}$  are shown by the dashed blue line.

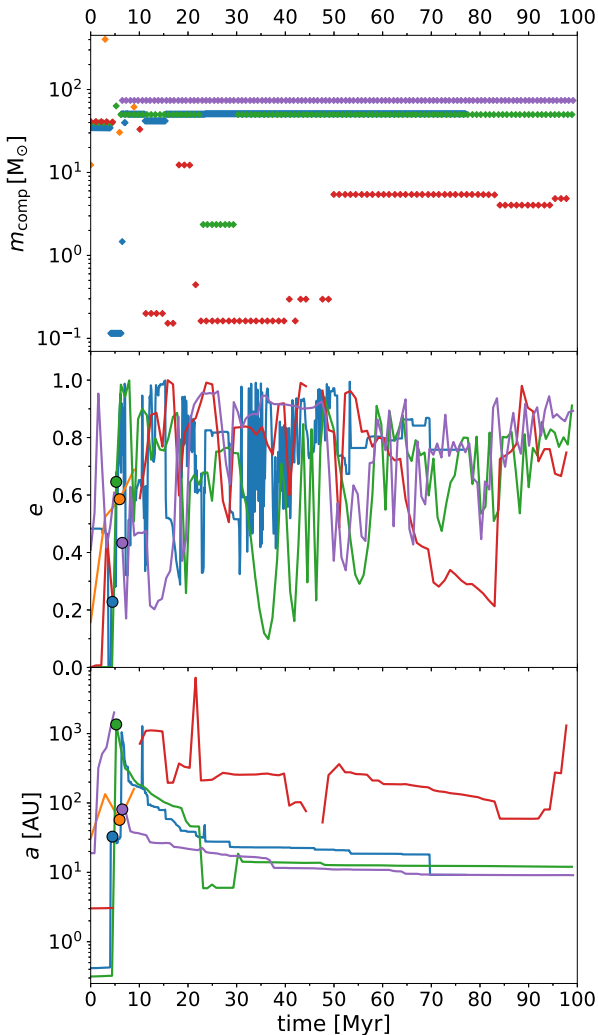
### 3.4 IMBHs in binary systems

IMBHs are the most massive objects which form in our simulations, and therefore they are likely to interact with other stars and form binary systems (Blecha et al. 2006; MacLeod, Trenti & Ramirez-Ruiz 2016). In our simulations, IMBHs spend  $\sim 85$  per cent of the time, on average, being part of a binary or triple system. Fig. 8 shows the evolution of the companion mass  $m_{\text{comp}}$ , orbital eccentricity  $e$ , and semimajor axis  $a$  for the five most massive IMBHs that form in our simulations. All these five massive IMBHs spend almost all their time in binary systems, except for one of them, which is ejected from the SC after  $\sim 10$  Myr.

These binaries are subject to continuous dynamical interactions which perturb their orbital properties. Eccentricity wildly oscillates and tends to assume values between 0.6 and 0.9 (Fig. 8). Each variation of eccentricity corresponds to a smaller variation of the semimajor axis, which tends to shrink over time. At the end of the simulation, two of the IMBHs shown in Fig. 8 are still members of a binary system with a companion BH. We have re-simulated the host SCs of these IMBHs for up to 1 Gyr to check for possible mergers, but none of these binary systems does merge.<sup>2</sup>

The simulated IMBHs tend to couple with other massive members of the SC, because dynamical exchanges generally lead to the build-up of more and more massive binaries (Hills & Fullerton 1980). Fig. 9 shows the distribution of the stellar types of the IMBH's companions. Besides main-sequence stars, which are the most common stellar type in the simulated SCs, the most common binary companions of IMBHs are BHs and neutron stars.

<sup>2</sup>NBODY6+ + GPU does not include a treatment for post-Newtonian terms. We cannot exclude the possibility that dynamical interactions combined with a treatment of post-Newtonian terms might allow the merger of some of our IMBH–BH binaries.

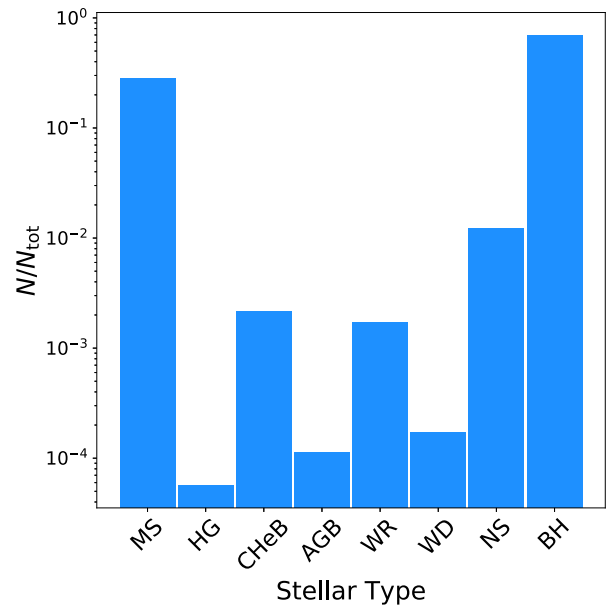


**Figure 8.** Evolution of the mass of the IMBH companion  $m_{\text{comp}}$  (upper panel), eccentricity  $e$  (middle panel), and semimajor axis  $a$  (lower panel) of the binaries which contain the five most massive IMBHs formed in our simulations. The coloured filled circles represent the formation time of the IMBH, i.e. the time when the progenitor star becomes an IMBH. The red circle is missing because the progenitor of this IMBH is single when it collapses to a BH.

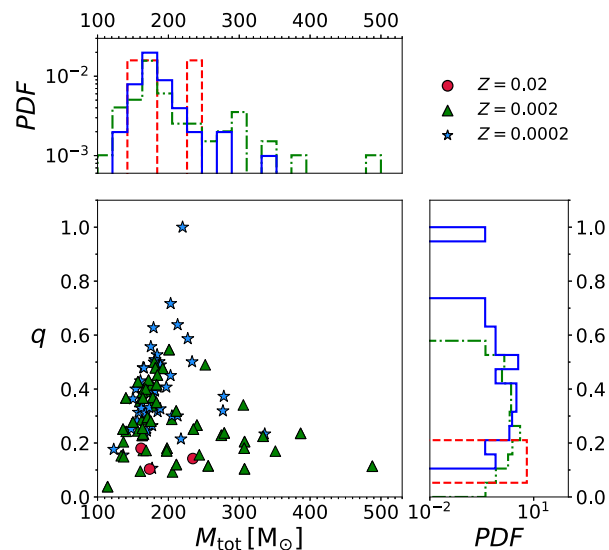
### 3.4.1 IMBH–BH binaries

We analyse the properties of all BBHs which host at least one IMBH in our simulations (hereafter, IMBH–BH binaries). From Fig. 9, it is apparent that IMBHs are extremely efficient in finding a companion BH:  $\sim 70$  per cent of all IMBHs reside in a BBH at the end of the simulations. Fig. 10 shows the total mass of the binaries  $M_{\text{tot}} = m_1 + m_2$  as a function of the mass ratio  $q = m_2/m_1$  (with  $m_2 \leq m_1$ ) of all the IMBH–BH systems. The binary with the smallest mass ratio has  $q \sim 0.04$  and secondary mass  $m_2 \sim 4.2 M_{\odot}$ . The most massive IMBH–BH binary has masses  $m_1 \sim 438 M_{\odot}$  and  $m_2 \sim 50 M_{\odot}$ . Values of the mass ratio  $q \lesssim 0.6$  are the most likely ones, but we find some binaries with  $q$  up to  $\sim 0.99$ . Overall, the IMBH–BH binaries have lower mass ratios than other BBHs (see fig. 5 of Di Carlo et al. 2020b, where we show that the vast majority of BBHs in young SCs have  $q \sim 0.9$ –1). This is expected, because SCs with more than one IMBH are very rare, and IMBHs can thus couple only with lower mass BHs.

We find no mergers of IMBH–BH binaries in our simulations. Adding post-Newtonian terms to our simulations might enhance

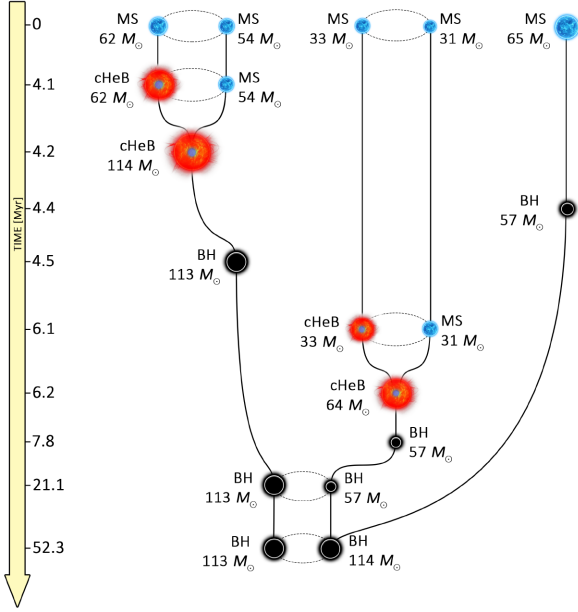


**Figure 9.** Distribution of the stellar types of the companions of all the simulated IMBHs throughout the entire simulation (0–100 Myr). The values on the y-axis are normalized to the total number of companions  $N_{\text{tot}}$ . The represented stellar types are: main sequence (MS); Hertzsprung gap (HG); core helium burning (CheB); asymptotic giant branch (AGB); Wolf-Rayet (WR); white dwarf (WD); neutron star (NS), and black hole (BH).



**Figure 10.** Mass ratio  $q = m_2/m_1$  versus total mass  $M_{\text{tot}} = m_1 + m_2$  of IMBH–BH binaries. Circles, triangles, and stars refer to  $Z = 0.02$ ,  $0.002$ , and  $0.0002$ , respectively. The marginal histograms show the distribution of  $q$  (on the y-axis) and  $M_{\text{tot}}$  (on the x-axis). The solid blue, dot-dashed green, and dashed red histograms refer to  $Z = 0.0002$ ,  $0.002$ , and  $0.02$ , respectively.

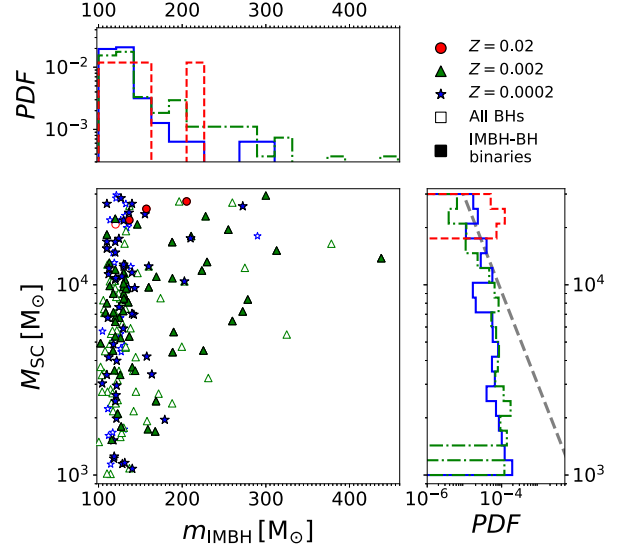
the chances of IMBH–BH merger during resonant encounters (e.g. Samsing 2018; Kremer et al. 2019; Zevin et al. 2019). However, our SCs are relatively small and disrupt quickly, so that these binaries do not have enough time to harden via dynamical interactions. Mergers of BBHs with at least one IMBH component seem to be extremely rare events even in much more massive SCs (e.g. Kremer et al. 2020).



**Figure 11.** Formation history of the binary IMBH found in our simulations. Main-sequence stars (with label MS) are represented as blue stars; core helium burning stars (label cHeB) are visualized as red stars with a blue core; black holes (label BH) are shown as black circles. The mass of each object is shown next to them. The time axis and the size of the objects are not to scale.

We find only one binary composed of two IMBHs (hereafter, binary IMBH), with masses  $m_1 \sim 113 M_\odot$  and  $m_2 \sim 114 M_\odot$ . Its formation history is shown in Fig. 11. First, two massive BHs of  $m_1 \sim 113 M_\odot$  and  $m_a \sim 57 M_\odot$ , which formed via runaway collision, dynamically couple and form a BBH. Then, a third BH of mass  $m_b \sim 57 M_\odot$  perturbs the BBH and in the process merges with the secondary BH of the binary, forming the binary IMBH. The formation of this binary involves both the runaway collision and the BBH merger mechanisms.

In our simulations, we do not include relativistic kicks. Were they included, they might have split this binary IMBH at birth. Thus, we calculated a posteriori the relativistic kick induced by the merger of two BHs with  $m_a = m_b = 57.4 M_\odot$ , according to the formulas presented in Maggiore (2018). To calculate the kick, we randomly drew the spin magnitudes of the two BHs from a Maxwellian distribution with  $\sigma = 0.1$  and the spin directions as isotropically distributed over the sphere (see e.g. Bouffanais et al. 2021 for a discussion of these assumptions). We repeated the calculation of the kick  $10^5$  times for different random draws of the spin magnitudes and tilts. We then calculated the probability that the binary system composed of the merger remnant and of the third BH with mass  $m_1 = 113 M_\odot$  remains bound after the relativistic kick. We calculate the mass of the merger remnant from Jiménez-Forteza et al. (2017), yielding  $m_{\text{rem}} \sim 109 M_\odot$ . To estimate the orbital properties of the binary IMBH with masses  $m_1$  and  $m_{\text{rem}}$  we start from an eccentricity  $e = 0.5$  and a semimajor axis  $a = 75.5$  au, which are the values given by NBODY6+ + GPU at the time of the merger between  $m_a$  and  $m_b$ , and we calculate the final semimajor axis and eccentricity (after the relativistic kick) following the equations detailed in the Appendix A of Hurley, Tout & Pols (2002). We find that the binary IMBH remains bound in  $\sim 38$  per cent of our random realizations. If the kicks are drawn from a Maxwellian distribution with  $\sigma = 0.01$  (very small kicks, as predicted by Fuller & Ma 2019), the binary IMBH remains bound in  $\sim 98$  per cent of the random realizations. The relativistic



**Figure 12.** Mass of the host SC ( $M_{\text{SC}}$ ) versus the mass of the IMBH ( $m_{\text{IMBH}}$ ). The marginal histograms show the distribution of  $M_{\text{SC}}$  (y-axis) and  $m_{\text{IMBH}}$  (x-axis). The filled symbols refer to IMBHs in IMBH–BH binaries, while the open symbols are single BHs. The solid blue, dot-dashed green, and dashed red colours refer to  $Z = 0.0002$ ,  $0.002$ , and  $0.02$ , respectively. The grey dashed line shows the initial mass function of SCs ( $dN/dM_{\text{SC}} \propto M_{\text{SC}}^{-2}$ ).

kick even favours the merger of the final binary IMBH in a few realizations ( $\sim 0.05$  per cent).

### 3.5 Properties of the host SCs

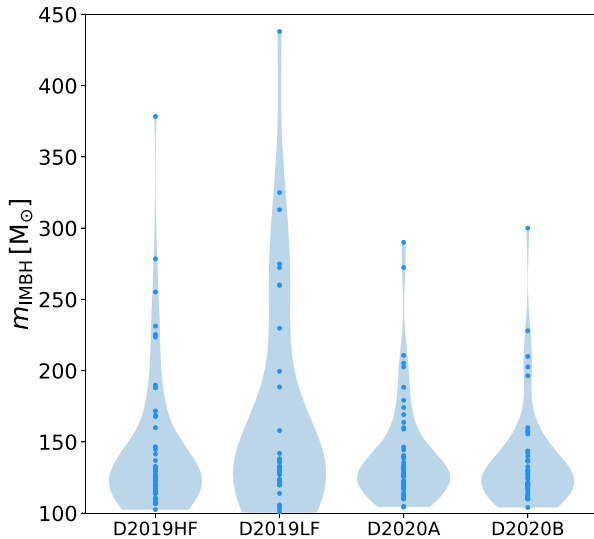
The presence of an IMBH may influence the evolution and the characteristics of the host SC (Mastrobuono-Battisti, Perets & Loeb 2014; Bianchini et al. 2015; Arca-Sedda 2016; Askar et al. 2017). Finding possible signatures of the presence of an IMBH in SCs may help us to understand whether observed SCs host or not an IMBH.

#### 3.5.1 SC mass

Fig. 12 shows the initial mass of the host SC as a function of the mass of the IMBH. The most massive IMBHs form preferentially in massive SCs. From the marginal histogram which shows the distribution of  $M_{\text{SC}}$ , it may seem that IMBHs are slightly more numerous in small SCs. However, since we adopted a mass function  $dN/dM_{\text{SC}} \propto M_{\text{SC}}^{-2}$ , we simulated many more small SCs than massive ones. As confirmed by Table 2, IMBHs form more efficiently in massive SCs where stellar collisions are more effective. For example, IMBHs form in  $\sim 8$  per cent of the SCs with  $10^4 M_\odot \leq M_{\text{SC}} \leq 3 \times 10^4 M_\odot$  and only in  $\lesssim 1$  per cent of the SCs with  $10^3 M_\odot \leq M_{\text{SC}} < 5 \times 10^3 M_\odot$ . In our simulations, IMBHs at solar metallicity form only in SCs with mass  $\gtrsim 20\,500 M_\odot$ . This feature may also be an effect of the fact that more massive SCs statistically tend to have more massive stars. For example, the median of the maximum initial stellar masses for SCs of mass  $10^3$ ,  $2.5 \times 10^3$ ,  $5 \times 10^3$ ,  $10^4$ , and  $3 \times 10^4 M_\odot$  is 54, 84, 108, 125, and  $140 M_\odot$ , respectively.

#### 3.5.2 Fractality and initial half-mass radius

In Fig. 13, we compare the IMBH mass distribution of different simulation sets. Simulations with low fractality (D2019LF) produce

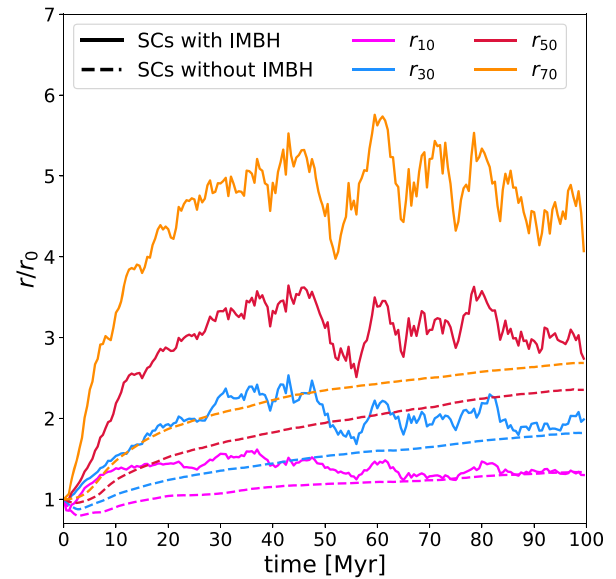


**Figure 13.** Distributions of IMBH masses for the different simulation sets. The blue filled circles mark the values of the masses, while the horizontal extent of each light blue region (violin plot) is proportional to the number of IMBHs at a given mass value. We refer to Table 1 for details on the different sets.

more massive IMBHs than high fractality ones (D2019HF). From Table 2, however, we see that the D2019HF set is more efficient than D2019LF in producing IMBHs. This means that high fractality helps to produce a larger number of IMBHs, but with lower mass. This happens because SCs with high fractality are composed of smaller and denser clumps (with a shorter two-body relaxation time-scale), where stellar collisions are more likely to occur. On the one hand, this increases their efficiency in forming IMBHs. On the other hand, a small clump does not host many massive stars (in our high-fractality SCs, a single clump may host up to two massive stars at most). Hence, in order to have a runaway collision of massive stars and to form a more massive IMBH, it is necessary that more clumps merge together before the death of their massive stars. In the high-fractality case, the various clumps merge together over a time-scale similar or longer than the IMBH formation time-scale. While it is easier to form a small IMBH in a single clump, stars from more clumps may be needed to form higher mass IMBHs.

In contrast, low fractality SCs are more similar to monolithic SCs and have a longer two-body relaxation time-scale. This means that the runaway collisions will begin later, involving all the massive stars of the SC. On the one hand, the mass of the collision product will not be limited by the mass of the clump, allowing the formation of more massive IMBHs. On the other hand, since the collisions start later, massive stars may die before assembling a sufficiently massive star. This makes low fractality SCs less efficient in forming IMBHs.

We now compare sets D2020A and D2020B in Fig. 13 and in Table 2. The SCs in D2020A have a smaller initial half-mass radius than in D2020B, but they both have the same degree of fractality ( $D = 1.6$ ). This means that the clumps will have approximately the same number of stars, but different densities. As shown in Fig. 1, the SCs in D2020A have a larger core density in the first  $\sim 8$  Myr, when the runaway collisions take place. Because of this higher initial density, set D2020A is more efficient in forming IMBHs, but the IMBH mass distributions are comparable.



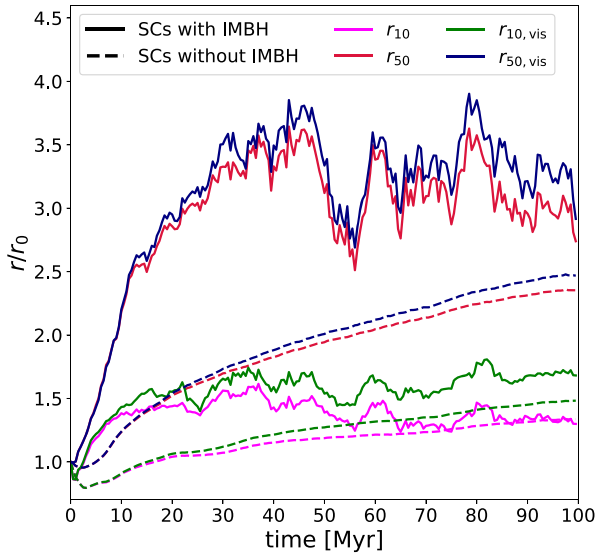
**Figure 14.** Time evolution of the median of the 10 per cent, 30 per cent, 50 per cent, and 70 per cent Lagrangian radii ( $r_{10}$ ,  $r_{30}$ ,  $r_{50}$ , and  $r_{70}$ ). The solid lines refer to SCs which contain at least one IMBH; dashed lines refer to SCs which do not contain IMBHs. Each radius is normalized to its initial value  $r_0$ . A simple moving average over five time-steps has been performed to filter out statistical fluctuations.

### 3.5.3 SC radii

We check if the presence of an IMBH affects the evolution of the Lagrangian radii of the simulated SCs. Fig. 14 shows the evolution of the 10 per cent, 30 per cent, 50 per cent, and 70 per cent Lagrangian radii of SCs with and without IMBHs. The SCs with IMBHs expand more rapidly in the first few Myrs than the SCs without IMBHs. Higher Lagrangian radii expand more, meaning that the expansion is stronger in the outer regions of the SC. This effect is a consequence of IMBH dynamics: the IMBH heats the SC, scattering stars to less bound orbits and making the cluster expand (Baumgardt, Makino & Ebisuzaki 2004b). After the initial expansion, the radii of the SCs with IMBHs flatten out. After  $\sim 65$  Myr, the 10 per cent Lagrangian radii ( $r_{10}$ ) of the SCs with and without IMBH overlap and start behaving in the same way. At the end of the simulations, the values of the 10 per cent and 30 per cent Lagrangian radii of SCs with and without IMBHs are almost the same. The presence of the IMBH has a stronger impact on Lagrangian radii in the first stages of the evolution of the SCs.

We also check the behaviour of Lagrangian radii calculated excluding neutron stars and BHs. In this way, we exclude the dark component and we obtain radii more similar to what observations can tell us, even if we do not account for any other observational bias (for example, we do not remove the lowest mass stars). Hereafter, we call this version of the Lagrangian radii *visible* Lagrangian radii. Fig. 15 shows the differences between the 10 per cent and 50 per cent Lagrangian radii and the visible 10 per cent and 50 per cent Lagrangian radii. Visible Lagrangian radii are slightly larger than the Lagrangian radii likely due to mass segregation, but the difference does not affect our main result: SCs with IMBHs tend to expand more than SCs without IMBHs. This result is consistent with previous studies (e.g. Baumgardt et al. 2004b; Baumgardt, Makino & Ebisuzaki 2004a; Trenti et al. 2007; Trenti, Vesperini & Pasquato 2010), which find that the presence of an IMBH in the SC leads to an initial stronger expansion.





**Figure 15.** Time evolution of the median of the 10 percent and 50 percent Lagrangian radii ( $r_{10}$  and  $r_{50}$ ) and of the corresponding *visible* Lagrangian radii  $r_{10, \text{vis}}$  and  $r_{50, \text{vis}}$  (i.e. the 10 percent and 50 percent Lagrangian radii calculated excluding BHs and neutron stars). The solid lines refer to SCs which contain at least one IMBH; dashed lines refer to SCs without IMBHs. Each radius is normalized to its initial value  $r_0$ . A simple moving average over five time-steps has been performed to filter out statistical fluctuations.

## 4 DISCUSSION

### 4.1 Accretion

In our simulations, we assume that if a star merges with a BH or with a neutron star the entire mass of the star is lost. Realistically, a fraction of the star’s mass could have accreted on to the compact object. Studies on tidal disruption events (e.g. Metzger & Stone 2016) show that this fraction is  $\sim 1$ –10 per cent. We also stress that most of the star–BH mergers in our simulations tend to involve low-mass stars, because massive stars turn into compact objects in a few Myr. Had we allowed for a higher accretion efficiency in our simulations, some of the formed IMBHs would have grown a slightly larger mass.

### 4.2 Relativistic kicks

We do not include relativistic kicks in our simulations. Were they included, what would have been their influence on our IMBH sample? The only IMBHs which can be affected by relativistic kicks are the nine IMBHs which form via BBH merger, which make up only 4 per cent of all the formed IMBHs. Eight of these IMBHs form via the merger of only one BBH. This means that all of them would have formed in any case, with the only difference that they could have been ejected from their host SC if the relativistic kick would have exceeded the SC’s escape velocity. Hence, we do not expect any significant difference in our results if we account for relativistic kicks.

### 4.3 Ejected IMBHs

We find that  $\sim 54$  per cent of all the formed IMBHs are ejected from the host SC, but what are the mechanisms behind such ejections? SC evaporation is not a major channel of IMBH ejection, because most IMBHs are ejected in the first tens of Myr of the life of their

parent SC. We find that 30 IMBHs escape while in binary systems and 88 IMBHs are ejected as single objects. To better understand the mechanisms behind the ejection of IMBHs in binary systems, we calculated the value of the semimajor axis below which an IMBH binary can be ejected by a single-binary scattering (Miller & Hamilton 2002; Antonini & Rasio 2016):

$$a_{\text{ej}} = \frac{\xi m_3^2}{(m_1 + m_2)^3} \frac{G m_1 m_2}{v_{\text{esc}}^2}, \quad (1)$$

where  $\xi = 3$  (Quinlan 1996) is a dimensionless parameter,  $v_{\text{esc}}$  is the escape velocity from the star cluster,  $m_1$  and  $m_2$  are the masses of the members of the binary, and  $m_3$  is the mass of the single particle. If we consider encounters with an object of mass 1, 5, 10  $M_{\odot}$ , we find  $a < a_{\text{ej}}$  for 0 per cent, 47 per cent, 93 per cent of the binaries. These results suggest that if single-binary encounters are the mechanism behind the ejection of IMBH–BH binaries, the single object is most likely a BH or a massive star. Furthermore, since  $a_{\text{ej}}$  depends on  $v_{\text{esc}}^{-2}$ , the semimajor axis threshold is larger in low-mass SCs, where wider binaries can thus escape.

## 5 CONCLUSIONS

IMBHs have mass in the range  $10^2$ – $10^5 M_{\odot}$  and bridge the gap between stellar-sized BHs and supermassive BHs. Currently, we lack unambiguous evidence of IMBH existence from electromagnetic observations, but we know several strong candidates (e.g. Farrell et al. 2009; Godet et al. 2014; Reines & Volonteri 2015; Kızıltan et al. 2017; Abbott et al. 2020a, b). Here, we have investigated the formation of IMBHs in young SCs through BBH mergers and the runaway collision mechanism.

In our simulations, 209 IMBHs form via the runaway collision mechanism, while only nine IMBHs form via BBH mergers. Each SC hosts a maximum of two IMBHs. We find IMBHs with mass up to  $\sim 438 M_{\odot}$ , but  $\sim 78$  per cent of all the simulated IMBHs have a mass between 100 and 150  $M_{\odot}$ : less massive IMBHs are more likely to form than the most massive ones.

As expected, IMBH formation is much less efficient at high metallicity, because stellar winds are more effective: only four IMBHs form at solar metallicity. The percentage of the simulated SCs which form at least one IMBH grows as metallicity decreases, going from 0.15 per cent at  $Z = 0.02$  to 3.5 per cent at  $Z = 0.0002$ .

IMBHs form more efficiently in massive SCs, where dynamics is more important. For example, IMBHs form in  $\sim 8$  per cent of the SCs with  $10^4 M_{\odot} \leq M_{\text{SC}} \leq 3 \times 10^4 M_{\odot}$  and in only  $\lesssim 1$  per cent of the SCs with  $10^3 M_{\odot} \leq M_{\text{SC}} < 5 \times 10^3 M_{\odot}$ .

In our simulations,  $\sim 54$  per cent of all the IMBHs are ejected from their parent SC, preferentially in the first  $\sim 25$  Myr from the beginning of the simulations. The IMBHs that remain inside the cluster rapidly sink towards the centre after their formation and stay there until the end of the simulation. IMBHs are more likely to be ejected if their mass is low ( $\lesssim 150 M_{\odot}$ ) and if the mass of their host SC is large ( $> 5 \times 10^3 M_{\odot}$ ).

Due to their high mass, IMBHs are likely to dynamically interact with other stars and form binary systems. In our simulations, IMBHs spend  $\sim 85$  per cent of the time, on average, being part of a binary or triple system. IMBHs tend to pair up with other massive BHs in the SC, because dynamical exchanges favour the formation of more massive binaries, which are more energetically stable (Hills & Fullerton 1980).

Overall, our simulations show that IMBHs form quite efficiently in massive young SCs via runaway stellar collisions, especially if the metallicity is relatively low ( $Z = 0.0002, 0.002$ ). While IMBHs

are efficient in pairing up dynamically, the occurrence of IMBH–BH mergers is extremely rare ( $<1$  every  $10^4$  young SCs).

## ACKNOWLEDGEMENTS

We thank the referee Mirek Giersz for his careful reading of the manuscript and for his many useful suggestions. MM, AB, UNDC, NG, GI, and SR acknowledge financial support from the European Research Council (ERC) under European Union’s Horizon 2020 research and innovation programme, Grant agreement no. 770017 (DEMOBLACK ERC Consolidator Grant). MP’s contribution to this material is based upon work supported by Tamkeen under the NYU Abu Dhabi Research Institute grant CAP3. NG acknowledges financial support by Leverhulme Trust Grant No. RPG-2019-350 and Royal Society Grant No. RGS-R2-202004.

## DATA AVAILABILITY

The data underlying this article will be shared on reasonable request to the corresponding authors.

## REFERENCES

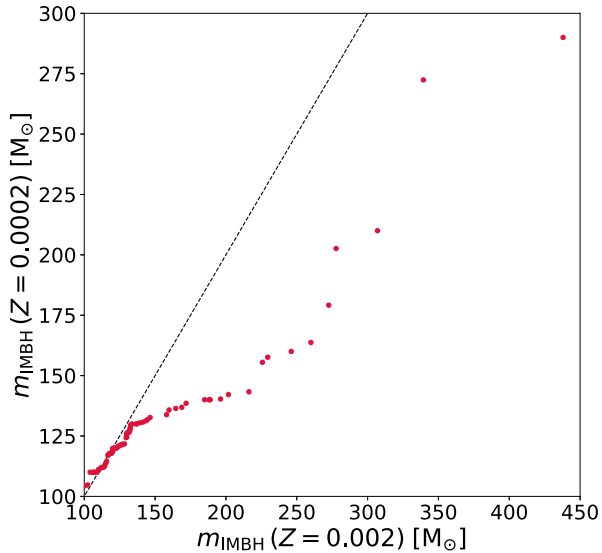
- Aasi J. et al., 2015, *Class. Quantum Gravity*, 32, 074001  
 Abbott R. et al., 2020a, *Phys. Rev. Lett.*, 125, 101102  
 Abbott R. et al., 2020b, *ApJ*, 900, L13  
 Acernese F. et al., 2015, *Class. Quantum Gravity*, 32, 024001  
 Amaro-Seoane P. et al., 2017, preprint (arXiv:1702.00786)  
 Anders E., Grevesse N., 1989, *Geochim. Cosmochim. Acta*, 53, 197  
 Anderson J., van der Marel R. P., 2010, *ApJ*, 710, 1032  
 Antonini F., Rasio F. A., 2016, *ApJ*, 831, 187  
 Antonini F., Gieles M., Gualandris A., 2019, *MNRAS*, 486, 5008  
 Arca-Sedda M., 2016, *MNRAS*, 455, 35  
 Arca Sedda M., Mapelli M., Spera M., Benacquista M., Giacobbo N., 2020, *ApJ*, 894, 133  
 Askar A., Bianchini P., de Vita R., Giersz M., Hypki A., Kamann S., 2017, *MNRAS*, 464, 3090  
 Baldassare V. F., Reines A. E., Gallo E., Greene J. E., 2015, *ApJ*, 809, L14  
 Barth A. J., Ho L. C., Rutledge R. E., Sargent W. L. W., 2004, *ApJ*, 607, 90  
 Baumgardt H., 2017, *MNRAS*, 464, 2174  
 Baumgardt H., Makino J., Ebisuzaki T., 2004a, *ApJ*, 613, 1133  
 Baumgardt H., Makino J., Ebisuzaki T., 2004b, *ApJ*, 613, 1143  
 Baumgardt H. et al., 2019, *MNRAS*, 488, 5340  
 Beccari G. et al., 2019, *ApJ*, 876, 87  
 Belloni D., Kroupa P., Rocha-Pinto H. J., Giersz M., 2018, *MNRAS*, 474, 3740  
 Bianchini P., Norris M. A., van de Ven G., Schinnerer E., 2015, *MNRAS*, 453, 365  
 Blecha L., Ivanova N., Kalogera V., Belczynski K., Fregeau J., Rasio F., 2006, *ApJ*, 642, 427  
 Bond J. R., Arnett W. D., Carr B. J., 1984, *ApJ*, 280, 825  
 Bouffanais Y., Mapelli M., Santoliquido F., Giacobbo N., Di Carlo U. N., Rastello S., Artale M. C., Iorio G., 2021, preprint (arXiv:2102.12495)  
 Carr B., Kühnel F., Sandstad M., 2016, *Phys. Rev. D*, 94, 083504  
 Chen Y., Bressan A., Girardi L., Marigo P., Kong X., Lanza A., 2015, *MNRAS*, 452, 1068  
 Chon S., Omukai K., 2020, *MNRAS*, 494, 2851  
 Clesse S., Garcia-Bellido J., 2020, preprint (arXiv:2007.06481)  
 Colgate S. A., 1967, *ApJ*, 150, 163  
 Das A., Schleicher D. R. G., Leigh N. W. C., Boehholt T. C. N., 2021, *MNRAS*, 503, 1051  
 De Luca V., Desjacques V., Franciolini G., Pani P., Riotto A., 2021, *Phys. Rev. Lett.*, 126, 051101  
 Di Carlo U. N., Giacobbo N., Mapelli M., Pasquato M., Spera M., Wang L., Haardt F., 2019, *MNRAS*, 487, 2947  
 Di Carlo U. N., Mapelli M., Bouffanais Y., Giacobbo N., Santoliquido F., Bressan A., Spera M., Haardt F., 2020a, *MNRAS*, 497, 1043  
 Di Carlo U. N. et al., 2020b, *MNRAS*, 498, 495  
 Di Cintio P., Pasquato M., Simon-Petit A., Yoon S.-J., 2021, preprint (arXiv:2103.02424)  
 Dong X.-B., Ho L. C., Yuan W., Wang T.-G., Fan X., Zhou H., Jiang N., 2012, *ApJ*, 755, 167  
 Fan X. et al., 2001, *AJ*, 122, 2833  
 Farrell S. A., Webb N. A., Barret D., Godet O., Rodrigues J. M., 2009, *Nature*, 460, 73  
 Favata M., Hughes S. A., Holz D. E., 2004, *ApJ*, 607, L5  
 Feng H., Soria R., 2011, *New Astron. Rev.*, 55, 166  
 Filippenko A. V., Ho L. C., 2003, *ApJ*, 588, L13  
 Filippenko A. V., Sargent W. L. W., 1989, *ApJ*, 342, L11  
 Fishbach M., Holz D. E., Farr B., 2017, *ApJ*, 840, L24  
 Fisher R. A., 1992, in *Breakthroughs in statistics*. Springer, New York, p. 66  
 Fragione G., Loeb A., Rasio F. A., 2020, *ApJ*, 902, L26  
 Freitag M., Rasio F. A., Baumgardt H., 2006a, *MNRAS*, 368, 121  
 Freitag M., Gürkan M. A., Rasio F. A., 2006b, *MNRAS*, 368, 141  
 Fryer C. L., Belczynski K., Wiktorowicz G., Dominik M., Kalogera V., Holz D. E., 2012, *ApJ*, 749, 91  
 Fujii M. S., Portegies Zwart S., 2013, *MNRAS*, 430, 1018  
 Fuller J., Ma L., 2019, *ApJ*, 881, L1  
 Gaburov E., Lombardi J. C., Portegies Zwart S., 2008, *MNRAS*, 383, L5  
 Gebhardt K., Rich R. M., Ho L. C., 2002, *ApJ*, 578, L41  
 Gebhardt K., Rich R. M., Ho L. C., 2005, *ApJ*, 634, 1093  
 Gerosa D., Berti E., 2017, *Phys. Rev. D*, 95, 124046  
 Gerosa D., Berti E., 2019, *Phys. Rev. D*, 100, 041301  
 Gerssen J., van der Marel R. P., Gebhardt K., Guhathakurta P., Peterson R. C., Pryor C., 2002, *AJ*, 124, 3270  
 Giacobbo N., Mapelli M., 2018, *MNRAS*, 480, 2011  
 Giacobbo N., Mapelli M., 2019, *MNRAS*, 482, 2234  
 Giacobbo N., Mapelli M., 2020, *ApJ*, 891, 141  
 Giacobbo N., Mapelli M., Spera M., 2018, *MNRAS*, 474, 2959  
 Giersz M., Leigh N., Hypki A., Lützgendorf N., Askar A., 2015, *MNRAS*, 454, 3150  
 Godet O., Lombardi J. C., Antonini F., Barret D., Webb N. A., Vingless J., Thomas M., 2014, *ApJ*, 793, 105  
 Goodman J., Hut P., 1989, *Nature*, 339, 40  
 Goodman J., Hut P., 1993, *ApJ*, 403, 271  
 Goodwin S. P., Whitworth A. P., 2004, *A&A*, 413, 929  
 Greene J. E., Ho L. C., 2004, *ApJ*, 610, 722  
 Greene J. E., Ho L. C., 2007, *ApJ*, 670, 92  
 Greene J. E., Strader J., Ho L. C., 2020, *ARA&A*, 58, 257  
 Gürkan M. A., Freitag M., Rasio F. A., 2004, *ApJ*, 604, 632  
 Heger A., Woosley S. E., 2002, *ApJ*, 567, 532  
 Hills J. G., Fullerton L. W., 1980, *AJ*, 85, 1281  
 Holley-Bockelmann K., Gültekin K., Shoemaker D., Yunes N., 2008, *ApJ*, 686, 829  
 Hurley J. R., Shara M. M., 2012, *MNRAS*, 425, 2872  
 Hurley J. R., Tout C. A., Pols O. R., 2002, *MNRAS*, 329, 897  
 Jiménez-Forteza X., Keitel D., Husa S., Hannam M., Khan S., Pürrer M., 2017, *Phys. Rev. D*, 95, 064024  
 Kaaret P., Prestwich A. H., Zezas A., Murray S. S., Kim D. W., Kilgard R. E., Schlegel E. M., Ward M. J., 2001, *MNRAS*, 321, L29  
 Kaaret P., Feng H., Roberts T. P., 2017, *ARA&A*, 55, 303  
 Kalogera V. et al., 2019, *BAAS*, 51, 242  
 Kawaguchi T., Kawasaki M., Takayama T., Yamaguchi M., Yokoyama J., 2008, *MNRAS*, 388, 1426  
 Kiziltan B., Baumgardt H., Loeb A., 2017, *Nature*, 542, 203  
 Kormendy J., Ho L. C., 2013, *ARA&A*, 51, 511  
 Kremer K. et al., 2019, *Phys. Rev. D*, 99, 063003  
 Kremer K. et al., 2020, *ApJ*, 903, 45  
 Kroupa P., 2001, *MNRAS*, 322, 231  
 Kroupa P., Boily C. M., 2002, *MNRAS*, 336, 1188  
 Küpper A. H. W., Maschberger T., Kroupa P., Baumgardt H., 2011, *MNRAS*, 417, 2300  
 Lada C. J., Lada E. A., 2003, *ARA&A*, 41, 57

- Lanzoni B. et al., 2013, *ApJ*, 769, 107
- Lin D. et al., 2018, *Nat. Astron.*, 2, 656
- Lin D. et al., 2020, *ApJ*, 892, L25
- Lousto C. O., Zlochower Y., 2011, *Phys. Rev. Lett.*, 107, 231102
- Lützgendorf N., Kissler-Patig M., Noyola E., Jalali B., de Zeeuw P. T., Gebhardt K., Baumgardt H., 2011, *A&A*, 533, A36
- Lützgendorf N. et al., 2013, *A&A*, 552, A49
- MacLeod M., Trenti M., Ramirez-Ruiz E., 2016, *ApJ*, 819, 70
- Madau P., Quataert E., 2004, *ApJ*, 606, L17
- Madau P., Rees M. J., 2001, *ApJ*, 551, L27
- Maggiore M., 2018, *Gravitational Waves: Volume 2: Astrophysics and Cosmology*, Gravitational Waves. Oxford Univ. Press, Oxford, available at <https://books.google.it/books?id=3ZNODwAAQBAJ>
- Maggiore M. et al., 2020, *J. Cosmol. Astropart. Phys.*, 2020, 050
- Mann C. R. et al., 2019, *ApJ*, 875, 1
- Mapelli M., 2016, *MNRAS*, 459, 3432
- Mapelli M., Ripamonti E., Zampieri L., Colpi M., Bressan A., 2010, *MNRAS*, 408, 234
- Mapelli M., Giacobbo N., Ripamonti E., Spera M., 2017, *MNRAS*, 472, 2422
- Mapelli M., Santoliquido F., Bouffanais Y., Arca Sedda M., Giacobbo N., Artale M. C., Ballone A., 2020a, preprint ([arXiv:2007.15022](https://arxiv.org/abs/2007.15022))
- Mapelli M., Spera M., Montanari E., Limongi M., Chieffi A., Giacobbo N., Bressan A., Bouffanais Y., 2020b, *ApJ*, 888, 76
- Mapelli M. et al., 2021, *MNRAS*, 505, 339
- Marks M., Kroupa P., 2012, *A&A*, 543, A8
- Mason D. M., Schuenemeyer J. H., 1983, *Ann. Stat.*, 11, 933
- Mastrobuono-Battisti A., Perets H. B., Loeb A., 2014, *ApJ*, 796, 40
- Matsumoto H., Tsuru T. G., Koyama K., Awaki H., Canizares C. R., Kawai N., Matsushita S., Kawabe R., 2001, *ApJ*, 547, L25
- McKernan B., Ford K. E. S., Lyra W., Perets H. B., 2012, *MNRAS*, 425, 460
- Merritt D., Milosavljević M., Favata M., Hughes S. A., Holz D. E., 2004, *ApJ*, 607, L9
- Metzger B. D., Stone N. C., 2016, *MNRAS*, 461, 948
- Mezcua M., 2017, *Int. J. Mod. Phys. D*, 26, 1730021
- Mezcua M., Roberts T. P., Lobanov A. P., Sutton A. D., 2015, *MNRAS*, 448, 1893
- Mezcua M., Civano F., Fabbiano G., Miyaji T., Marchesi S., 2016, *ApJ*, 817, 20
- Mezcua M., Civano F., Marchesi S., Suh H., Fabbiano G., Volonteri M., 2018, *MNRAS*, 478, 2576
- Miller M. C., Colbert E. J. M., 2004, *Int. J. Mod. Phys. D*, 13, 1
- Miller M. C., Hamilton D. P., 2002, *MNRAS*, 330, 232
- Miller-Jones J. C. A. et al., 2012, *ApJ*, 755, L1
- Moe M., Di Stefano R., 2017, *ApJS*, 230, 15
- Nyland K., Marvil J., Wrobel J. M., Young L. M., Zauderer B. A., 2012, *ApJ*, 753, 103
- Perera B. B. P. et al., 2017, *MNRAS*, 468, 2114
- Peters P. C., 1964, *Phys. Rev.*, 136, 1224
- Portegies Zwart S. F., McMillan S. L. W., 2002, *ApJ*, 576, 899
- Portegies Zwart S. F., Baumgardt H., Hut P., Makino J., McMillan S. L. W., 2004, *Nature*, 428, 724
- Portegies Zwart S. F., McMillan S. L. W., Gieles M., 2010, *ARA&A*, 48, 431
- Punturo M. et al., 2010, *Class. Quantum Gravity*, 27, 194002
- Quinlan G. D., 1996, *New Astron.*, 1, 35
- Raccanelli A., Kovetz E. D., Bird S., Cholis I., Muñoz J. B., 2016, *Phys. Rev. D*, 94, 023516
- Reines A. E., Volonteri M., 2015, *ApJ*, 813, 82
- Reines A. E., Condon J. J., Darling J., Greene J. E., 2020, *ApJ*, 888, 36
- Reitze D. et al., 2019, *BAAS*, 51, 35
- Rizzuto F. P. et al., 2021, *MNRAS*, 501, 5257
- Rodriguez C. L., Zevin M., Amaro-Seoane P., Chatterjee S., Kremer K., Rasio F. A., Ye C. S., 2019, *Phys. Rev. D*, 100, 043027
- Samsing J., 2018, *Phys. Rev. D*, 97, 103014
- Sana H. et al., 2012, *Science*, 337, 444
- Sanders R. H., 1970, *ApJ*, 162, 791
- Sasaki M., Suyama T., Tanaka T., Yokoyama S., 2016, *Phys. Rev. Lett.*, 117, 061101
- Scelfo G., Bellomo N., Raccanelli A., Matarrese S., Verde L., 2018, *J. Cosmol. Astropart. Phys.*, 2018, 039
- Schneider R., Ferrara A., Natarajan P., Omukai K., 2002, *ApJ*, 571, 30
- Seth A. C. et al., 2010, *ApJ*, 714, 713
- Spera M., Mapelli M., 2017, *MNRAS*, 470, 4739
- Spitzer L., 1987, *Dynamical Evolution of Globular Clusters*. Princeton Univ. Press, Princeton
- Strader J., Chomiuk L., Maccarone T. J., Miller-Jones J. C. A., Seth A. C., Heinke C. O., Sivakoff G. R., 2012, *ApJ*, 750, L27
- Strohmayer T. E., Mushotzky R. F., 2003, *ApJ*, 586, L61
- Sugimoto D., Bettwieser E., 1983, *MNRAS*, 204, 19P
- Sutton A. D., Roberts T. P., Walton D. J., Gladstone J. C., Scott A. E., 2012, *MNRAS*, 423, 1154
- Swartz D. A., Ghosh K. K., Tennant A. F., Wu K., 2004, *ApJS*, 154, 519
- Tagawa H., Haiman Z., Bartos I., Kocsis B., Omukai K., 2021a, preprint ([arXiv:2104.09510](https://arxiv.org/abs/2104.09510))
- Tagawa H., Kocsis B., Haiman Z., Bartos I., Omukai K., Samsing J., 2021b, *ApJ*, 908, 194
- Takekawa S., Oka T., Iwata Y., Tsujimoto S., Nomura M., 2019, *ApJ*, 871, L1
- Takekawa S., Oka T., Iwata Y., Tsujimoto S., Nomura M., 2020, *ApJ*, 890, 167
- Tanaka K., 2018, *ApJ*, 859, 86
- Tanikawa A., Susa H., Yoshida T., Trani A. A., Kinugawa T., 2021, *ApJ*, 910, 30
- Tremou E. et al., 2018, *ApJ*, 862, 16
- Trenti M., Ardi E., Mineshige S., Hut P., 2007, *MNRAS*, 374, 857
- Trenti M., Vesperini E., Pasquato M., 2010, *ApJ*, 708, 1598
- van der Marel R. P., 2004, in Ho L. C., ed., *Coevolution of Black Holes and Galaxies*. Cambridge University Press, Cambridge, p. 37
- van der Marel R. P., Anderson J., 2010, *ApJ*, 710, 1063
- Vink J. S., Muijres L. E., Anthonisse B., de Koter A., Gräfenor G., Langer N., 2011, *A&A*, 531, A132
- Volonteri M., 2010, *A&AR*, 18, 279
- Wang L., Spurzem R., Aarseth S., Nitadori K., Berczik P., Kouwenhoven M. B. N., Naab T., 2015, *MNRAS*, 450, 4070
- Wang L. et al., 2016, *MNRAS*, 458, 1450
- Zevin M., Samsing J., Rodriguez C., Haster C.-J., Ramirez-Ruiz E., 2019, *ApJ*, 871, 91
- Zocchi A., Gieles M., Hénault-Brunet V., 2019, *MNRAS*, 482, 4713

## APPENDIX A: COMPARING THE HIGH-MASS TAILS OF THE IMBH MASS DISTRIBUTIONS AT DIFFERENT METALLICITIES

Here, we explore if the mass of the heaviest IMBHs formed in our simulations depends on their progenitors' metallicity. In particular, Fig. 2 seems to suggest that the most massive IMBHs form at intermediate metallicity ( $Z = 0.002 \sim 0.1 Z_{\odot}$ ) rather than at the lowest considered metallicity ( $Z = 0.0002 \sim 0.01 Z_{\odot}$ ). To assess this trend with a more quantitative evaluation, we compare the end tail of the IMBH mass distribution at  $Z = 0.0002$  and  $Z = 0.002$ . The mass distribution of all the BHs in the  $Z = 0.002$  and the  $Z = 0.0002$  samples are significantly different, with  $p = 2.2 \times 10^{-16}$  according to a Kolmogorov–Smirnov (K-S) test. This rises to  $p = 0.08$  when only IMBHs (with mass  $M > 100 M_{\odot}$ ) are considered, likely because the K-S test is notoriously insensitive to distribution tails and to small samples (see e.g. Mason & Schuenemeyer 1983).

Fig. A1 shows the quantiles of the IMBH masses for  $Z = 0.0002$  against those with  $Z = 0.002$ . This quantile–quantile plot allows us to visualize how the two distributions differ: samples drawn from the same distribution would result in points straddling the identity line (black dashed line in Fig. A1), as each and every quantile would be the same for both samples, modulo random fluctuations. This is indeed the case for our two samples in Fig. A1 up until about  $130 M_{\odot}$ ,



**Figure A1.** Quantile–quantile plot for the IMBH masses in SCs with  $Z = 0.002$  and  $Z = 0.0002$ . The  $x$ -axis shows the values corresponding to selected quantiles of the  $Z = 0.002$  sample, while the values of the  $y$ -axis represent the same quantiles for  $Z = 0.0002$ . If the IMBH masses were extracted from the same distribution, deviations from the diagonal 1:1 black dashed line would be due to random fluctuations only.

but for higher IMBH masses the  $Z = 0.002$  simulations appear to produce IMBHs that are systematically heavier than the  $Z = 0.0002$  simulations. For example, the 90th percentile is  $228 M_{\odot}$  for the  $Z = 0.002$  and  $157 M_{\odot}$  for  $Z = 0.0002$ .

It is still far from trivial to conclude whether the high-mass tails and in particular the maximum mass produced is different. Comparing maxima is an intrinsically hard problem, because, even for well-behaved distributions, the variance of extremes is much larger than that of e.g. the mean. A widespread non-parametric approach for testing differences in maxima relies on an exact Fisher test (Fisher 1992) applied to the contingency table obtained by counting how many BHs lie above selected mass cut-offs in each of the two metallicity samples. This is similar to applying a  $\chi^2$  test, but is exact for the case in which cell counts are very low, as we expect them to be for large enough mass cut-offs. We find that the null hypothesis (that both metallicities produce IMBHs above and below each mass cut-off in the same proportion) cannot be rejected for cut-offs at the 90th, 95th, and 99th percentiles of the joined samples, so we conclude that at the moment our data do not constrain the impact of metallicity on maximum IMBH mass, at least if we forego making assumptions on the underlying parent distribution.

A general concern with this comparison could be that the number of simulations run at the two different metallicities is different, so the samples can be compared fairly only through a bootstrap resampling that takes this difference into account. We thus repeated our analysis by randomly resampling from the  $Z = 0.002$  sample either a number of IMBHs equal to that in the less numerous  $Z = 0.0002$  sample or a number of IMBHs corresponding to an equivalent number of simulations. Even with this test the results from the above still hold, i.e. no significant difference in the high-mass tails of the distributions could be ascertained non-parametrically.

This paper has been typeset from a  $\text{\TeX}/\text{\LaTeX}$  file prepared by the author.